Observational Properties of Weakly Coupled DM

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Talk based on arXiv: 1506.04048, 1601.07733, 1604.02401

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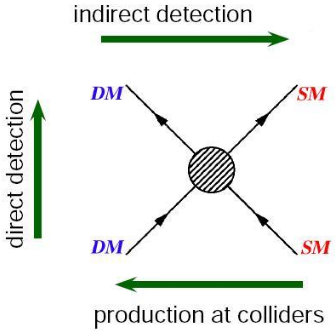
What is Dark Matter?

- What is the correct explanation for the invisible matter content observed in the universe? Does the dark matter particle exist? Or are there many dark matter particles?
- Are they WIMP's, FIMP's, SIMP's, GIMP's, PIDM's, WISP's, ALP's, Wimpzillas, or sterile neutrinos? Or should gravity be modified?

Image: Chandra X-ray Observatory

Search for Dark Matter

- Many on-going experiments exist to find the correct explanation



- But... what if dark matter interacts only very feebly with the known particles, or not at all?

Original image: Max-Planck-Institut Für Kernphysik

The Model

- The scalar sector of the model is specified by the potential

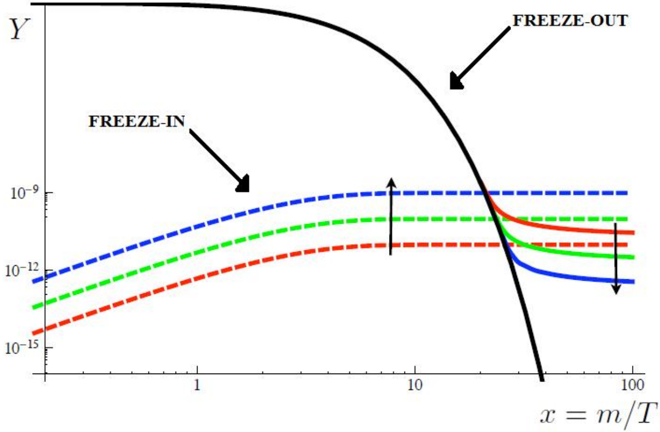
$$V(\Phi, s) = \mu_{\rm h}^2 \Phi^\dagger \Phi + \lambda_{\rm h}(\Phi^\dagger \Phi)^2 + \frac{1}{2}\mu_{\rm s}^2 s^2 + \frac{\lambda_{\rm s}}{4} s^4 + \frac{\lambda_{\rm sh}}{2} \Phi^\dagger \Phi s^2$$

- Here Φ and s are, respectively, the usual Standard Model Higgs doublet and a real singlet scalar.

- The coupling between Φ and s acts as a portal between the Standard Model and an unknown Dark Sector (the so-called Higgs portal).

Dark Matter production mechanisms

- There are basically two mechanisms for dark matter production: freeze-out and freeze-in



The original image is from Hall et al. (arXiv:0911.1120)

The Freeze-Out

- Dark matter is initially in thermal equilibrium with the SM particles. This requires a rather strong coupling, $\lambda_{\rm sh} \simeq 0.1$.
- May lead to a WIMP miracle: thermal relic with weak cross-section and a mass $m_{\rm s} \sim$ EW scale gives the right relic abundance.
- Starts to be very constrained by experiments[1]

[1] For a recent review, see e.g. M. Klasen, M. Pohl, G. Sigl (arXiv: 1507.03800)

Frozen-in Dark Matter

- Requires $\lambda_{\rm sh} \lesssim 10^{-7}$, or otherwise the singlet sector thermalizes with the SM (this is sometimes called a FIMP scenario)
- Is produced from many different sources including thermal bath of Standard Model particles and primordial scalar condensates[2]
- Leaves observable imprints on CMB[3]
- Cannot be tested in colliders but can be tested by cosmological and astrophysical observations[4]

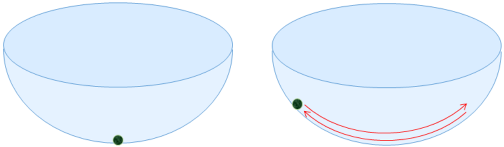
[2] S. Nurmi, TT, K. Tuominen (arXiv: 1506.04048)

[3] K. Kainulainen, S. Nurmi, TT, K. Tuominen, V. Vaskonen (arXiv: 1601.07733)

[4] M. Heikinheimo, TT, K. Tuominen, V. Vaskonen (arXiv:1604.02401)

Dark Matter from a primordial field

- During cosmic inflation, scalar fields typically acquire fluctuations proportional to the inflationary scale[5], $h, s \simeq H_* \lesssim 10^{14}$ GeV



Scalar fields fluctuate during cosmic inflation.

- After inflation, these scalar condensates will decay to particles

- The end products can constitute Dark Matter

[5] Starobinsky & Yokoyama (arXiv:astro-ph/9407016)

Imprints on CMB

- The observational bounds are significantly different depending on whether the singlet constitutes isocurvature or adiabatic dark matter
- The dark matter component sourced by a primordial scalar field is isocurvature and therefore strictly constrained by CMB observations[6]:

$$\frac{\Omega_{\rm DM} h^2}{0.12} \lesssim 10^{-5} \lambda_s^{-1/4} \tag{1}$$

[6] See K. Kainulainen, S. Nurmi, TT, K. Tuominen, V. Vaskonen (arXiv: 1601.07733) for details

The isocurvature bound

- To constrain the dark sector parameters, we compute $\Omega_{\rm DM} = \Omega_{\rm DM}(\lambda_{\rm s}, m_{\rm DM}, H_*)$ from theory[7]

$$\frac{\Omega_{\rm DM} h^2}{0.12} \simeq 10^{-4} \lambda_{\rm s}^{-5/8} \left(\frac{m_{\rm DM}}{\rm GeV}\right) \left(\frac{H_*}{10^{11}\rm GeV}\right)^{3/2}, \tag{2}$$

  and combine it with the isocurvature bound, $\Omega_{\rm DM} h^2/0.12 \lesssim 10^{-5} \lambda_{\rm s}^{-1/4}$.

- For fixed $m_{\rm DM}, H_*$, this gives a lower bound on $\lambda_{\rm s}$

- Note: $\Omega_{\rm DM}$ depends on H_* $\Rightarrow$ a novel connection between the dark matter abundance and the inflationary scale

[7] See S. Nurmi, TT, K. Tuominen (arXiv: 1506.04048) and K. Kainulainen, S. Nurmi, TT, K. Tuominen, V. Vaskonen (arXiv: 1601.07733) for details

Dark Matter self-interactions

- Astrophysical observations provide an upper bound on DM self-interactions[8]

$$\frac{\sigma_{\rm DM}}{m_{\rm DM}} = \frac{9\lambda_{\rm s}^2}{32\pi m_{\rm s}^3} \lesssim 1\,\frac{\rm cm^2}{\rm g} \tag{3}$$

- What kind of constraints do these limits place together?

[8] See e.g. D. Harvey et al. (arXiv: 1503.07675)

Thermal history of the Dark Sector

- An initial population of DM is produced through Higgs decays $h \to ss$ at $T \sim m_{\rm h}$. In the standard freeze-in scenario, this is also the final abundance.
- However, if the number changing interactions $2 \to 4$ in the dark sector are fast, they will lead to thermalization of the dark sector
- This reduces the average momentum of DM particles and increases their number density until thermal equilibrium is reached

Dark Freeze-out

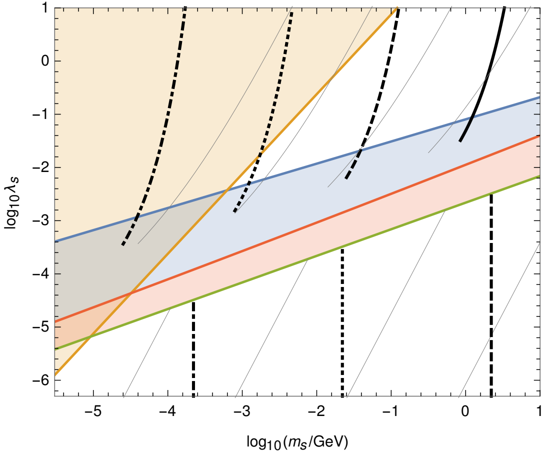
- The $2 \leftrightarrow 4$ interactions maintain thermal equilibrium until the $4 \rightarrow 2$ interaction rate drops below the Hubble rate and the number density freezes out

- This mechanism is referred to as dark freeze-out

- By knowing the initial DM abundance sourced by Higgs decays, the resulting DM relic density can be computed from the conservation of entropy

The results

- Three regimes: The dark freeze-out (above red line), the standard freeze-in (below green line), no solution at all (red)
- Two constraints: DM self-interactions (yellow), isocurvature perturbations (gray contours for different H_*'s)



Conclusions

- Thermal history of dark sector contains many interesting features, which have been studied only vaguely
- Cosmological and astrophysical observations provide a valuable resource on testing different dark matter models
- We have derived stringent constraints on Higgs portal dark matter model and found a novel connection between dark matter abundance and inflationary energy scale